

UHE neutrino follow-up of Gravitational Wave events with the Pierre Auger Observatory

Michael Schimp on behalf of the Pierre Auger Collaboration

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**PIERRE
AUGER**
OBSERVATORY



bmb+f

Großgeräte
der physikalischen
Grundlagenforschung



BERGISCHE
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WUPPERTAL

LIGO GW Events (GW150914 & GW151226)

- Binary BH mergers @ $d_L \sim 400$ Mpc
- $E_{\text{GW}} > M_{\odot} c^2$ ($\sim 10^{54}$ erg)



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Kotera and Silk arXiv:1602.06961

BH merger: possible GRB cause

- Needed: B field + debris
- Could produce the measured UHECR flux up to 100 EeV!
→ Needs $\sim 3\%$ efficiency ($E_{\text{UHECR}}/E_{\text{GW}}$)

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- **Fermi GBM detection 0.4 s** after GW150914, compatible direction
- **No prompt PeV ν** after GW150914 (IceCube, ANTARES)
- UHE (EeV) ν emission predicted (*Vietri, Waxman, Murase*)

Pierre Auger Observatory surface detector: **Large acceptance**

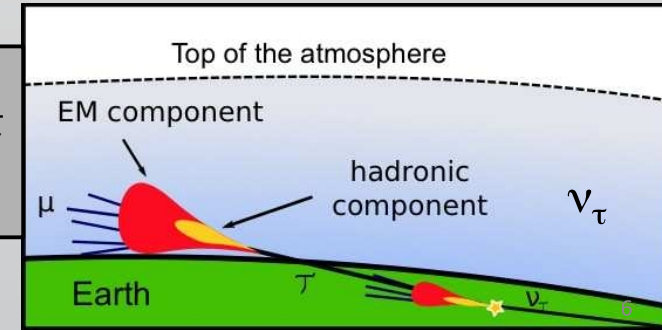


→ Well suited for UHE ν search!

Pierre Auger Observatory surface detector: Large acceptance



Earth-skimming ν_τ
($90^\circ < \theta < 95^\circ$)

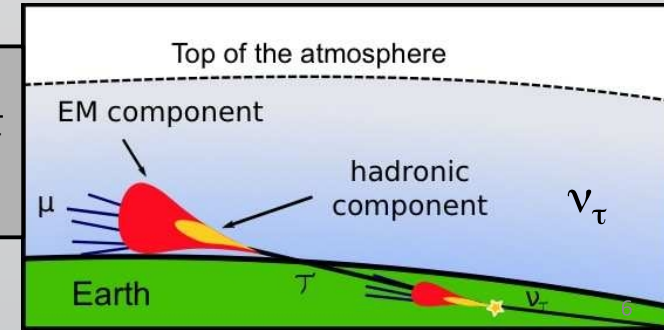


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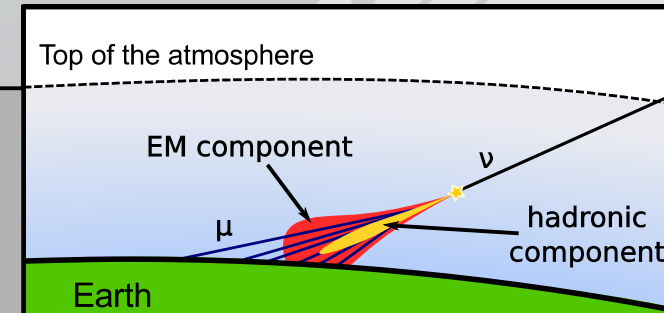
Pierre Auger Observatory surface detector: Large acceptance



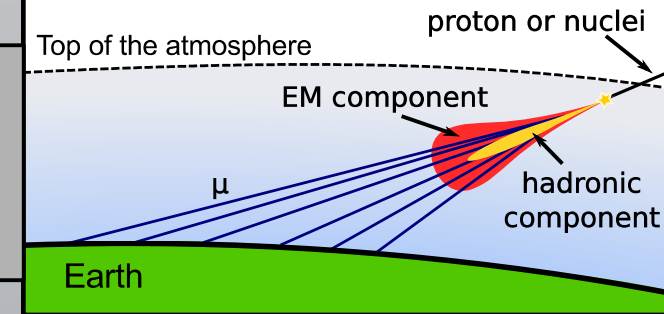
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Down-going ν
(signal)
($75^\circ < \theta < 90^\circ$)



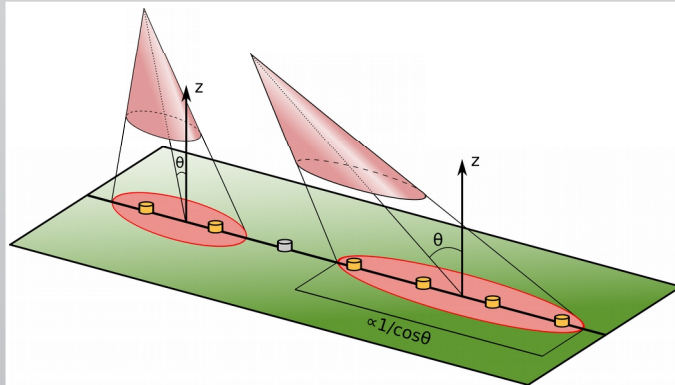
Down-going nuclei (bg)
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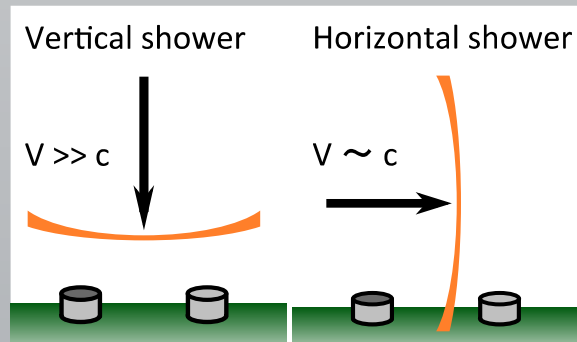
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UHE Neutrinos In The Pierre Auger Observatory

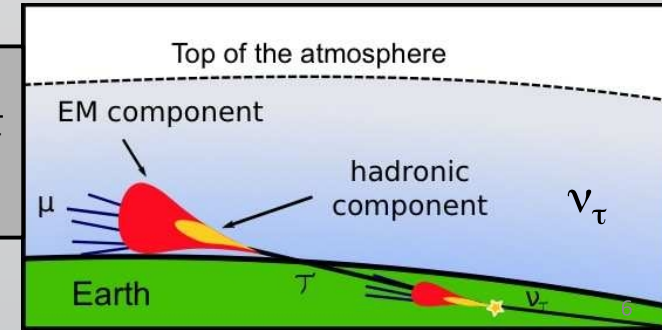
- Inclination: $\theta > 75^\circ$
- Elongated footprint



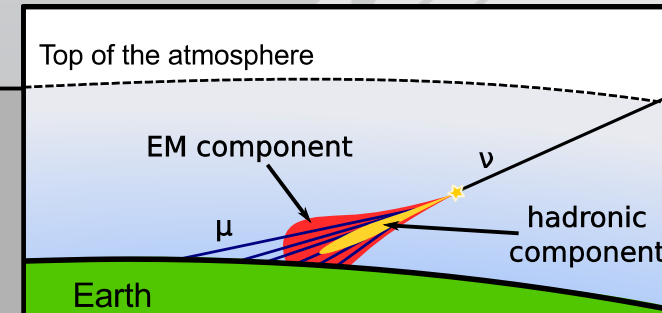
- “Ground signal speed” $\sim c/\cos(\theta)$



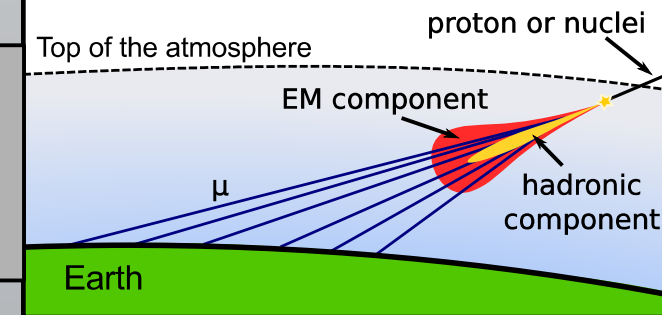
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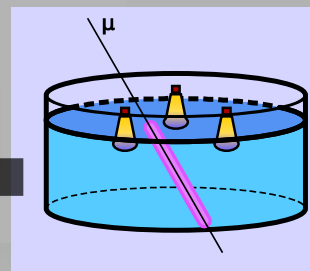
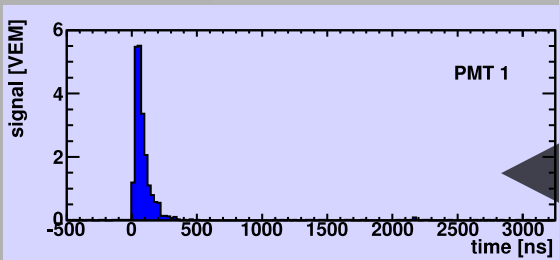
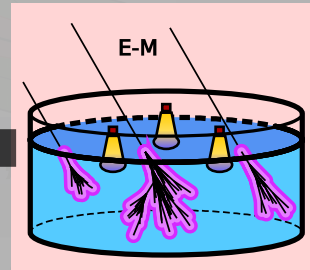
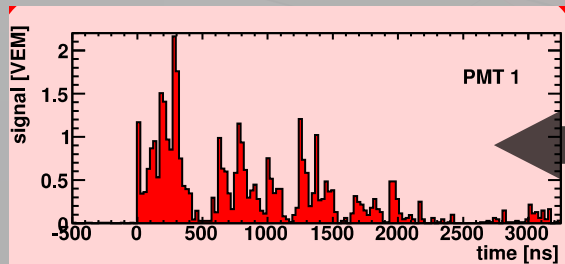
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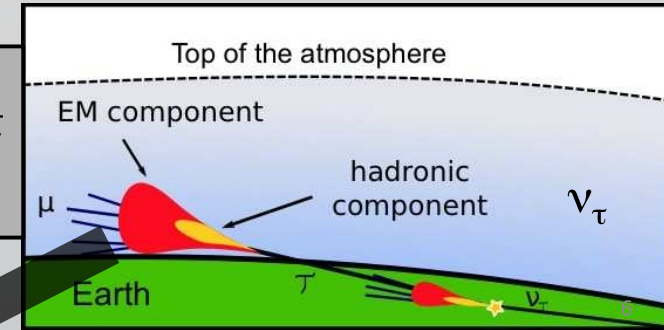
UHE Neutrinos In The Pierre Auger Observatory

Reject “muonic” events:

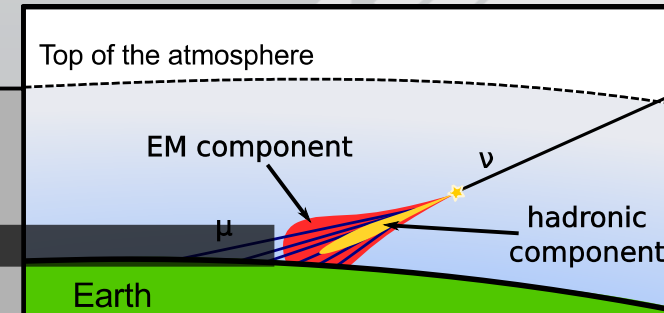
- Select events dominated by wide PMT signals
- Compare to simulation → ν identification



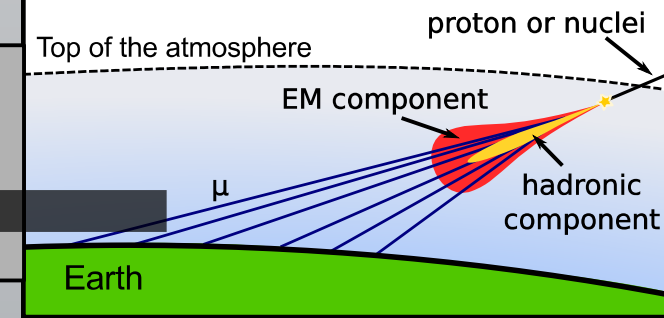
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No candidate in $[-500 \text{ s}, 1 \text{ day}]$
around GW events

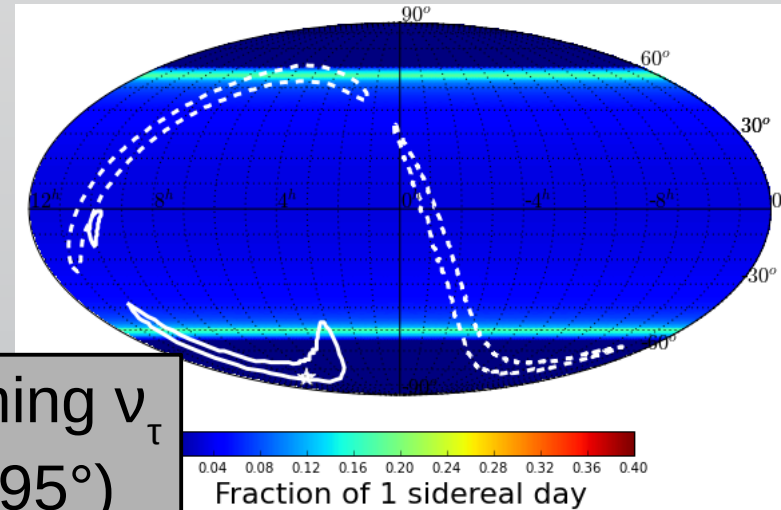
“Prompt” / “afterglow”

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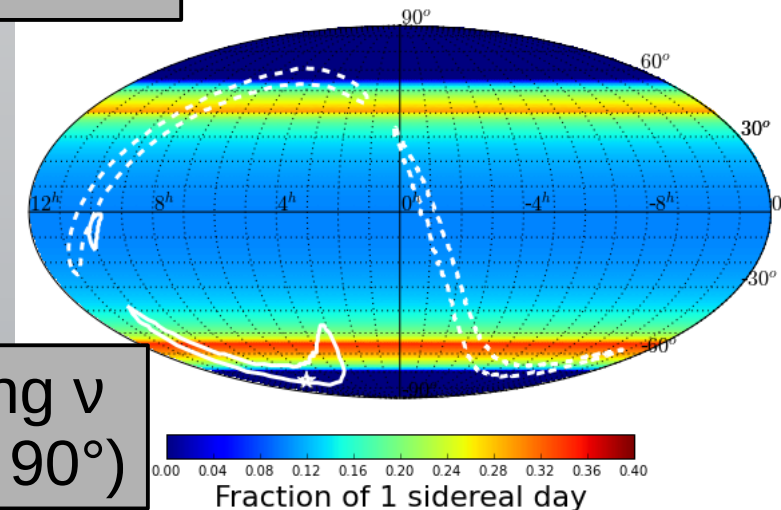
→ Calculate **exposure** taking
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- Time-dependent aperture
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- ν -nucleon cross section +
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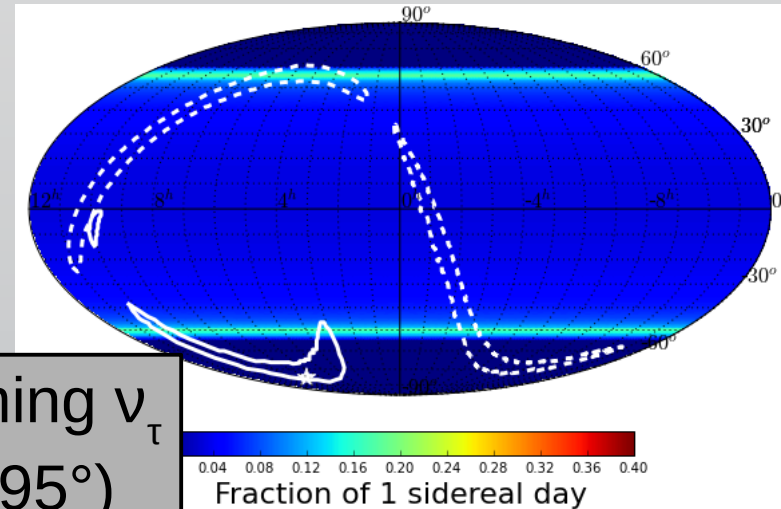
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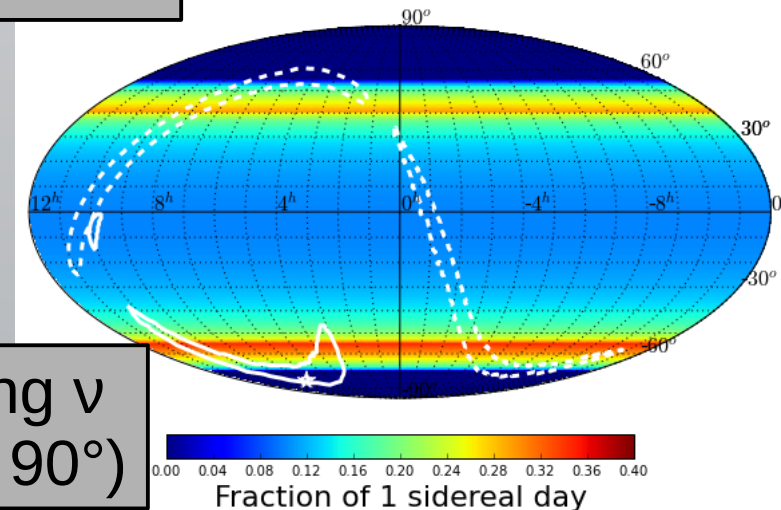
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→ Calculate upper limits on **energy radiated in UHE ν** (δ) (E^{-2} spectrum)

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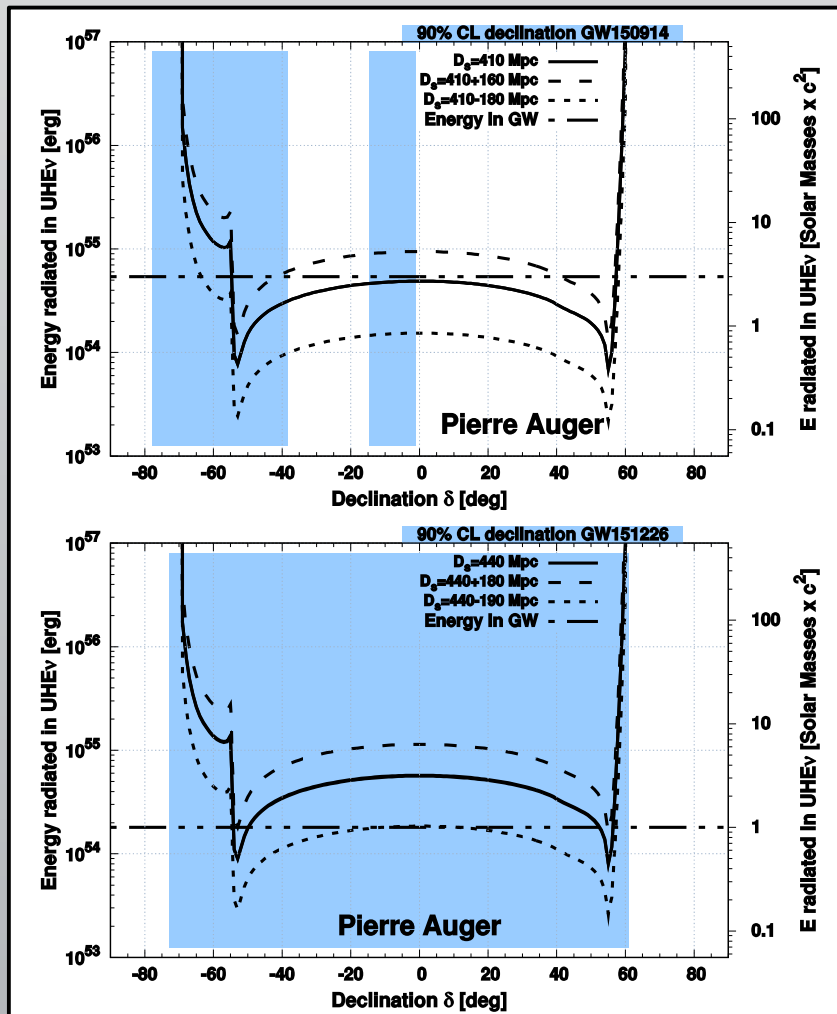


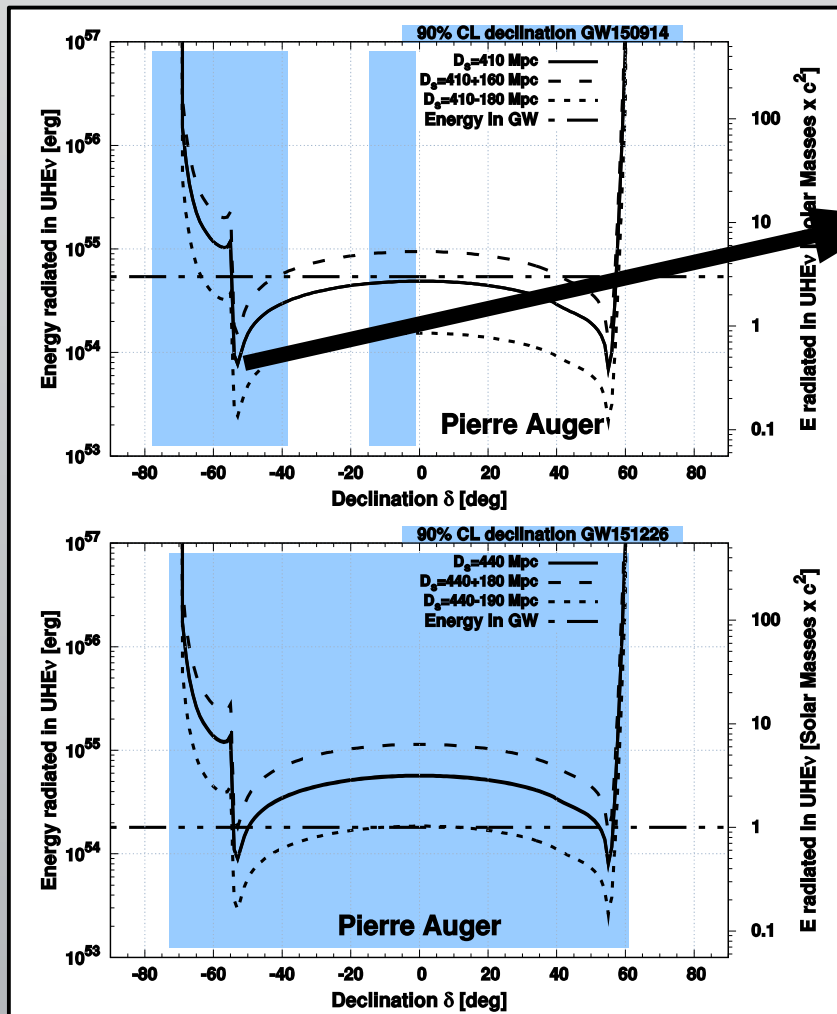
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Results & Conclusions

GW Follow-Up Results





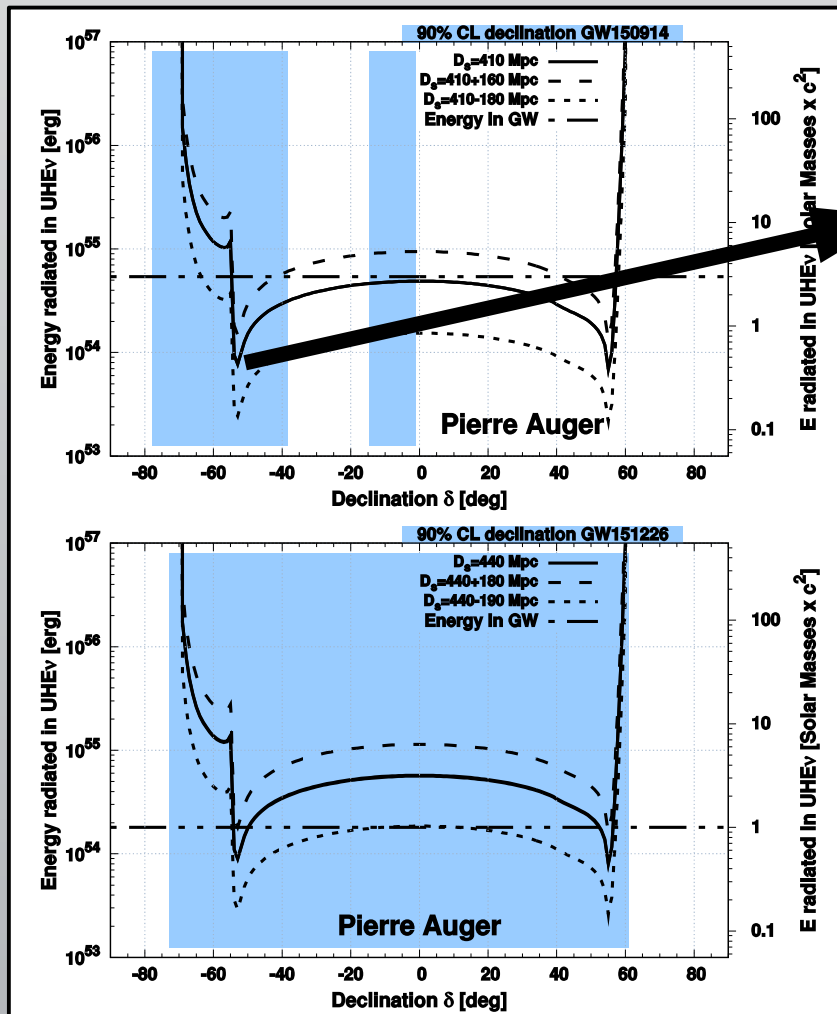
Most stringent:
15% of GW energy into UHEv

→ *Kotera, Silk*: 3% of GW energy needed for UHECR to explain their flux

- If 3% of GW energy into UHEv
→ ~ 1 event in our search (E^{-2})

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Also *Kotera, Silk*:

Prediction of **diffuse** UHEv flux from inferred BH merger **population** with 3% UHECR production efficiency:

- $E^2 dN/dE \sim 8.3 \times 10^{-8} \text{ GeV cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1} \times f_\nu \times f_z$
- Auger data up to 2013:
 $E^2 dN/dE < 6.4 \times 10^{-9} \text{ GeV cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$

ν optical depth

Redshift loss / source evolution

At least 1 is true

- UHECR production efficiency < 3%
- Inefficient charged pion (\rightarrow neutrino) production, $f_\nu < 1$
- Unfavorable source evolution

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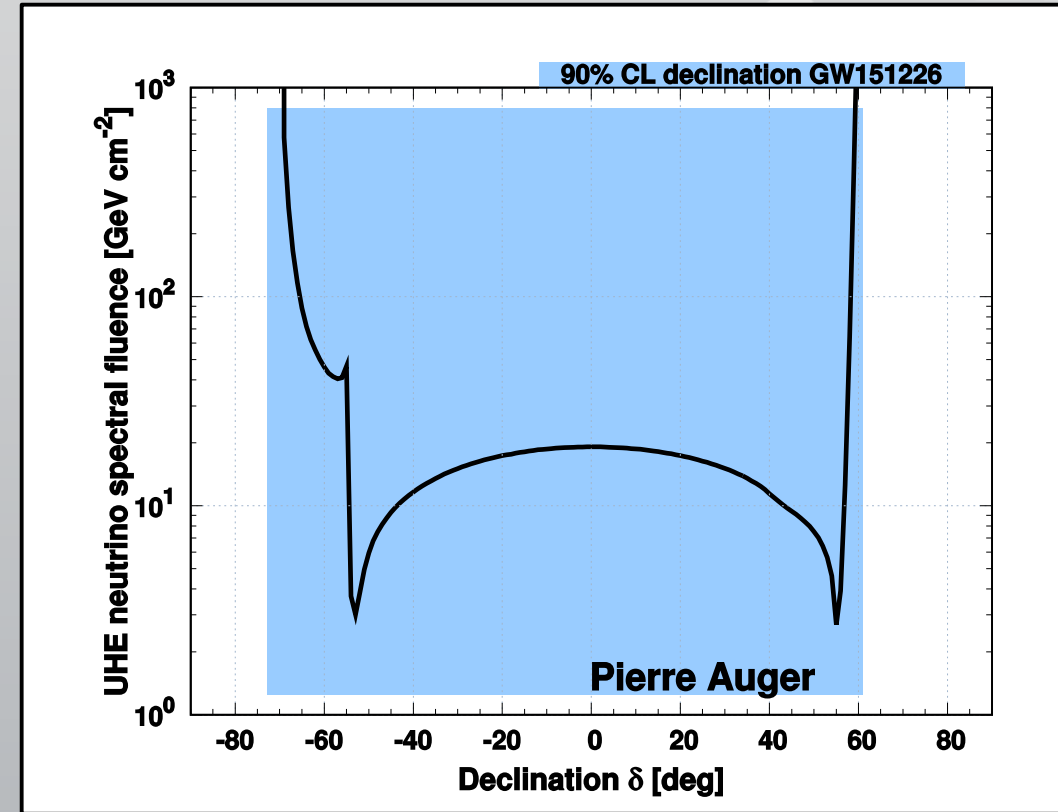
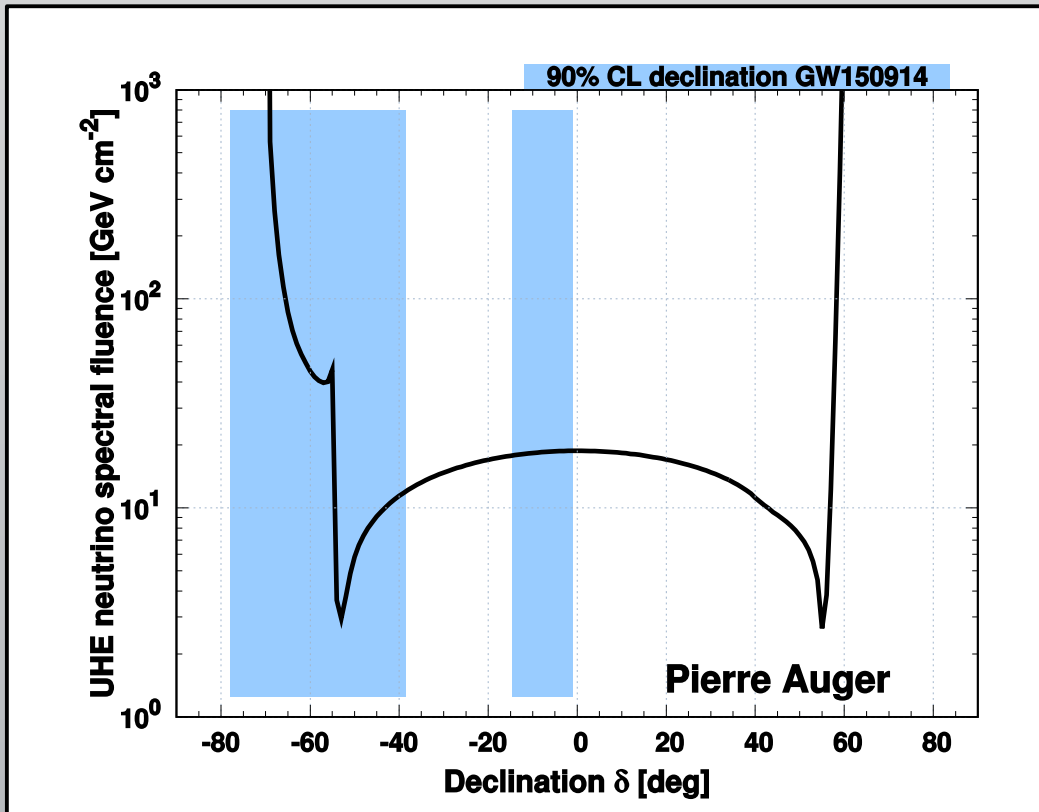
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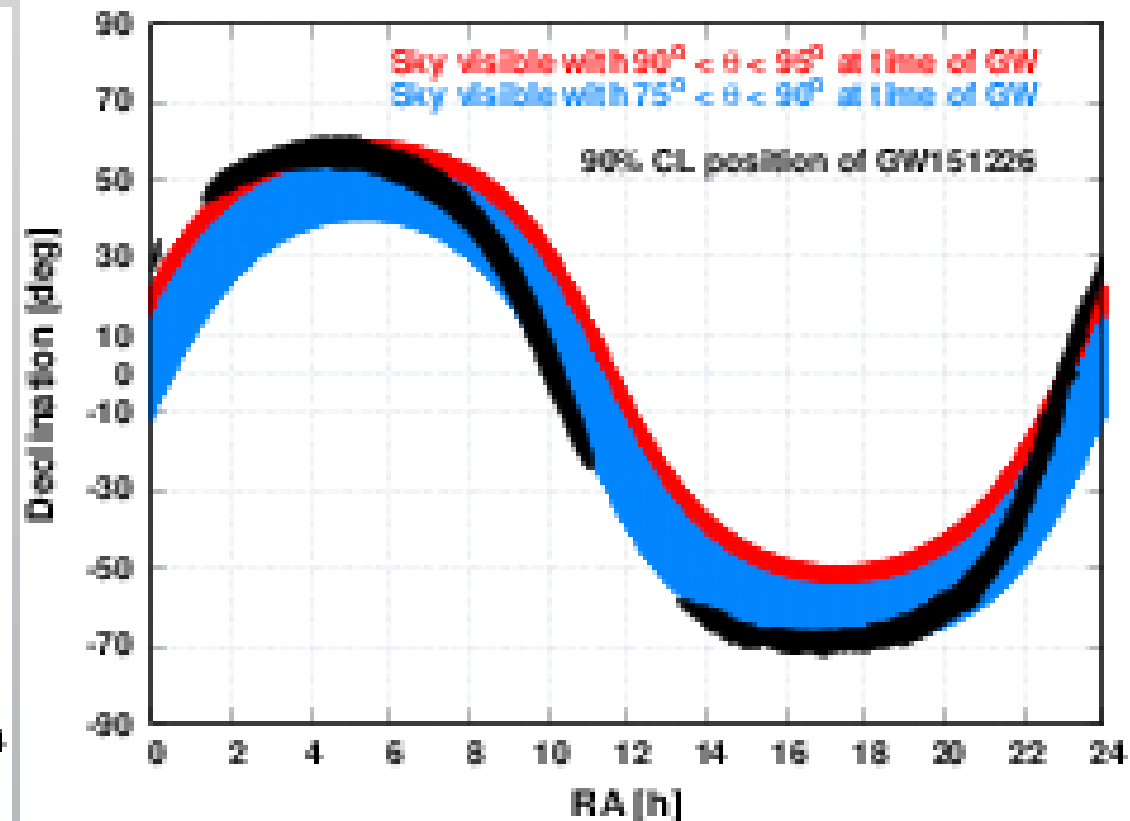
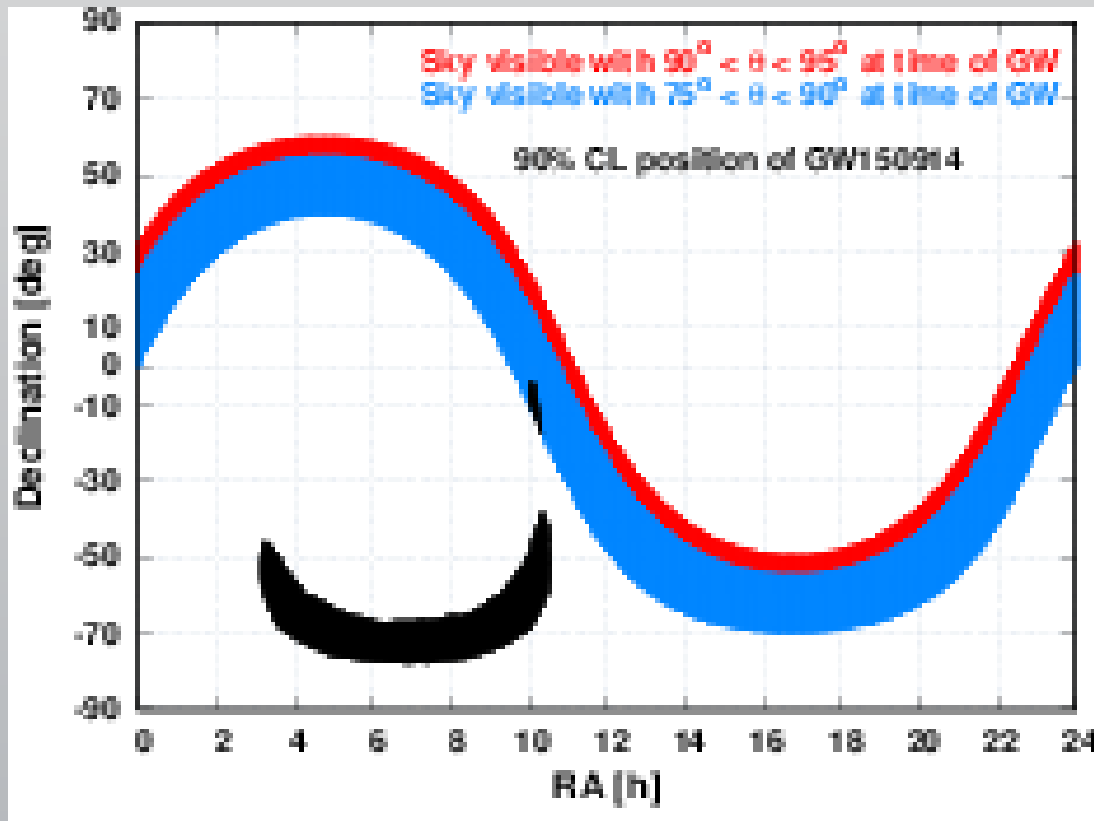
Extrapolation based
 \rightarrow We need **more** GW events!

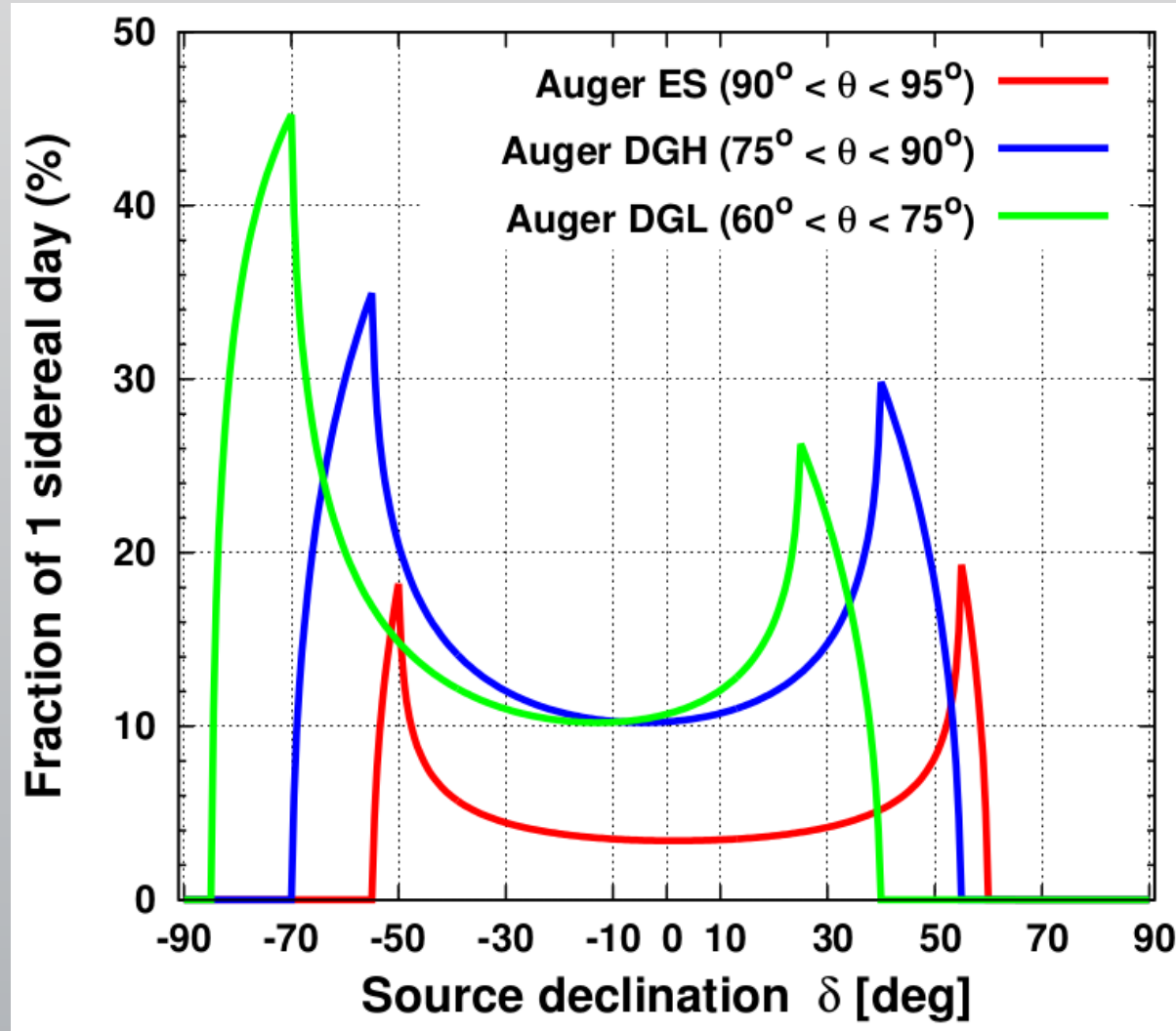
~~The End~~
—Questions?

Fluence Upper Limits

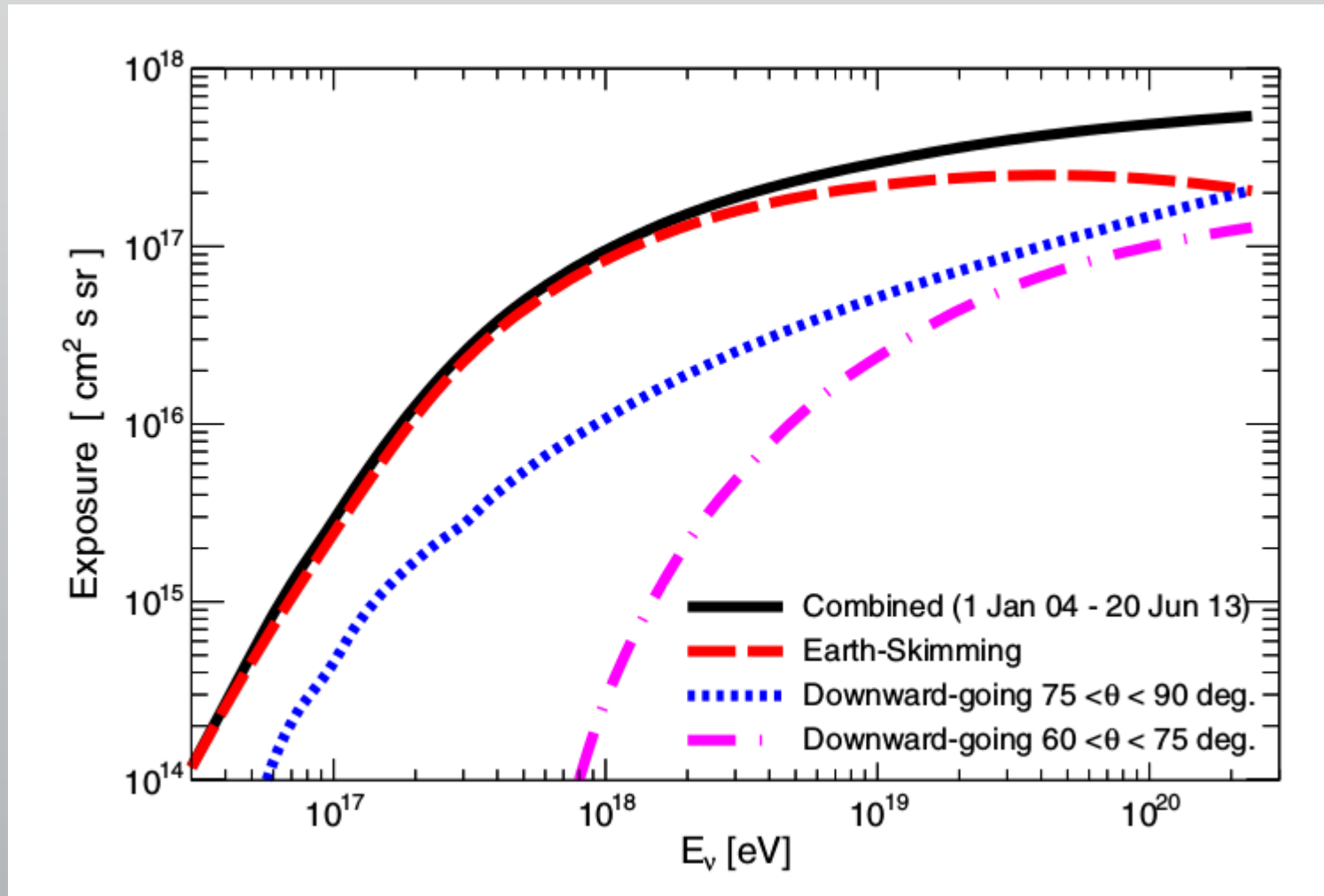


Visible Solid Angles At Times Of Events

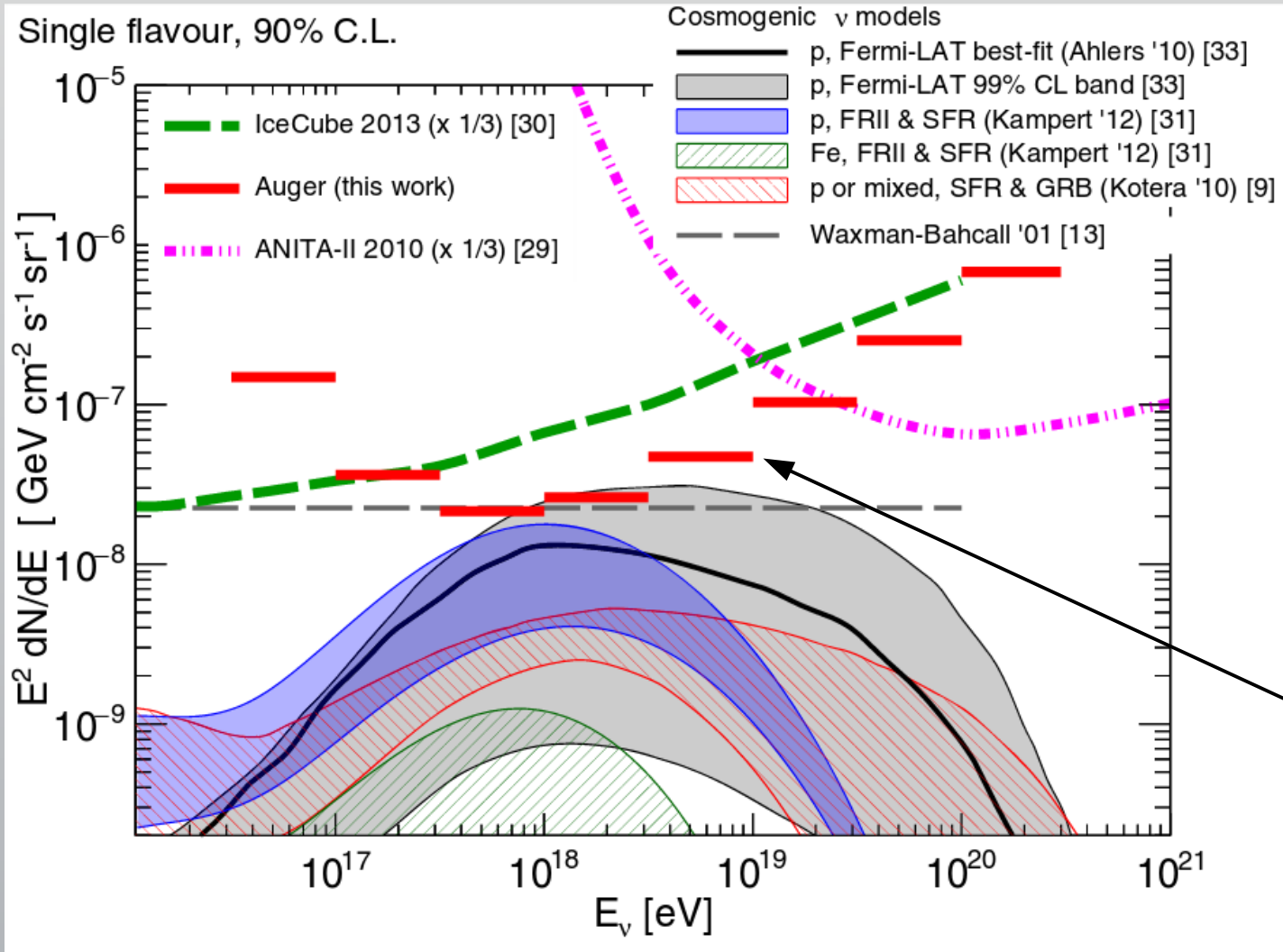




Energy Dependence Of Diffuse Exposure (Phys Rev D 91, 092008 (2015))

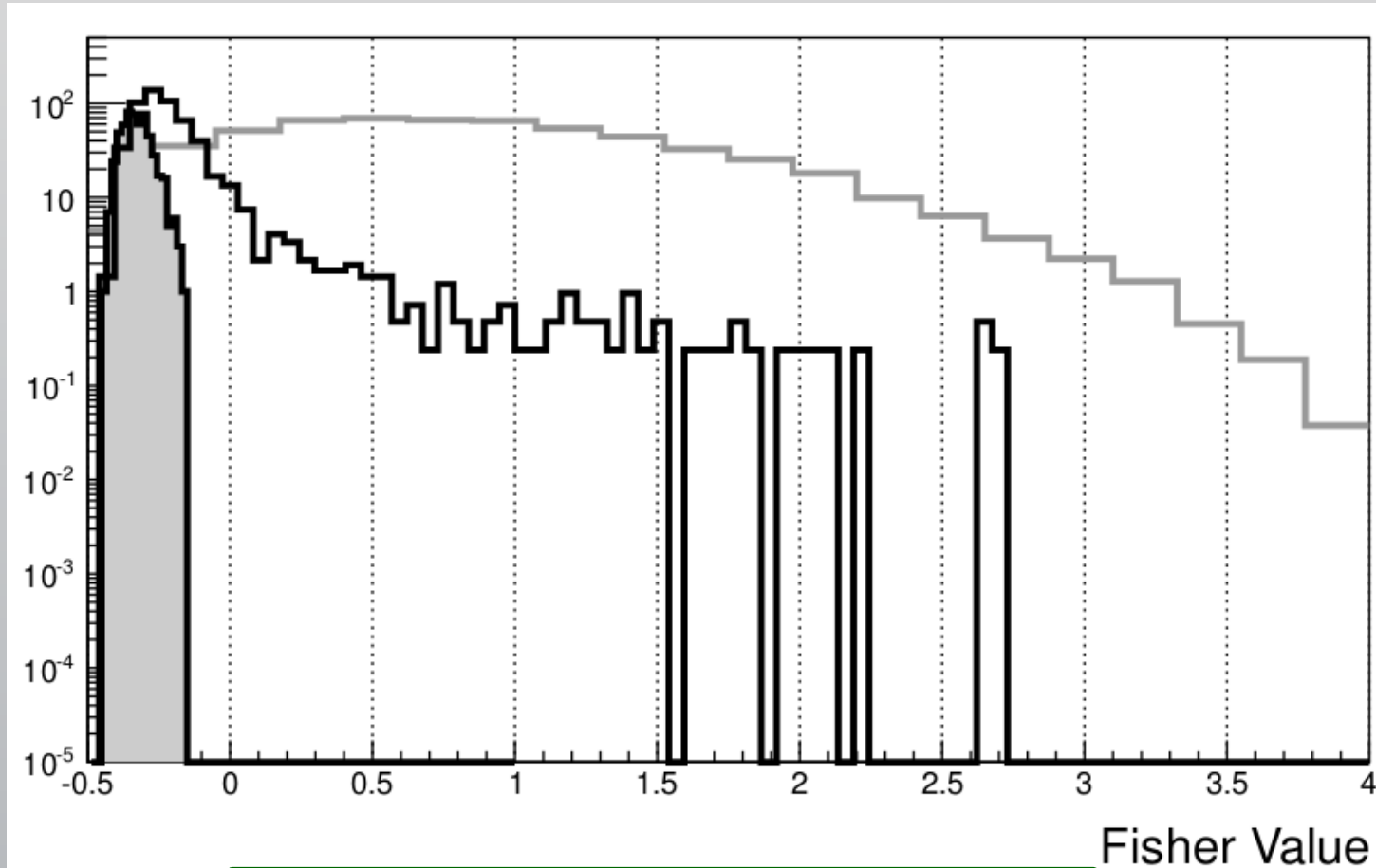


Binned Diffuse Flux Limits (Phys Rev D 91, 092008 (2015))

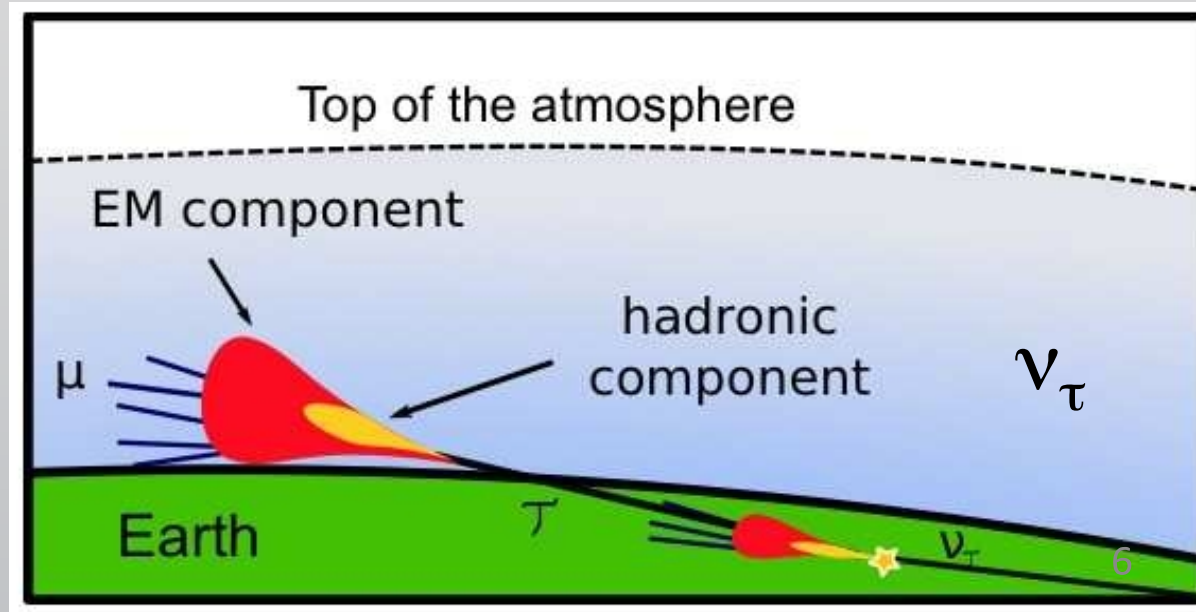


$$k = \frac{N_{\text{up}}}{\int_{E_\nu} E_\nu^{-2} \mathcal{E}_{\text{tot}}(E_\nu) dE_\nu}$$

Bin-wise
(0.5 in log E)



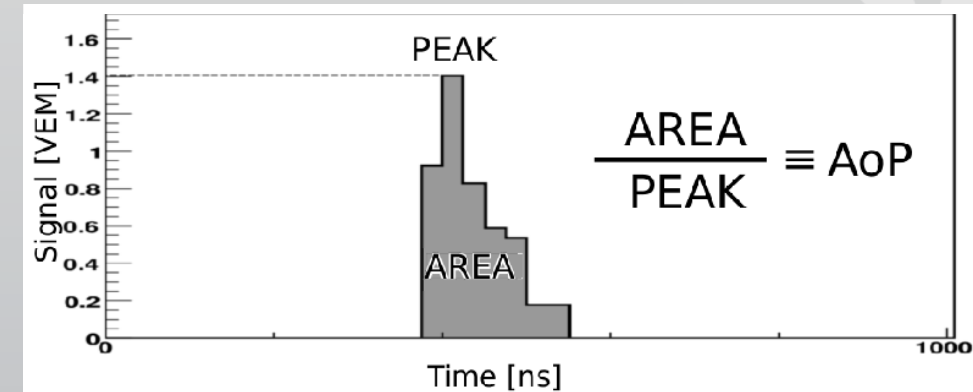
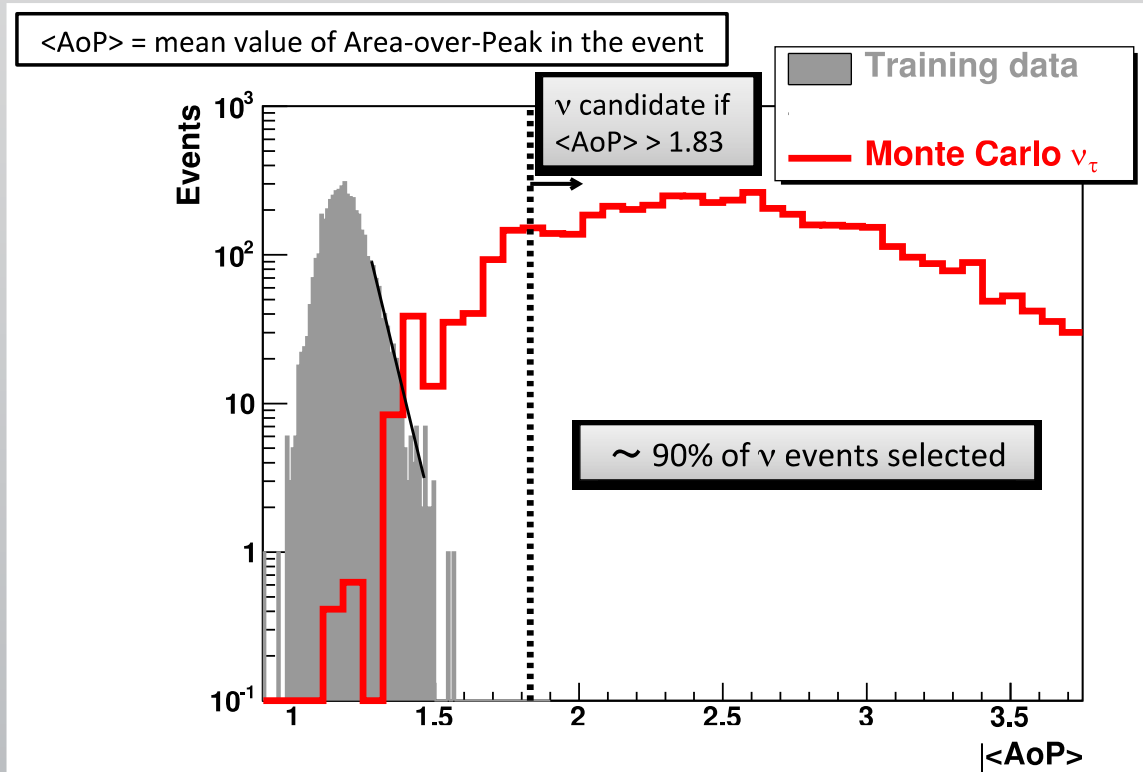
$$(d) 67.5^\circ < \theta_{\text{Rec}} \leq 70.5^\circ$$



Earth skimming events dominated by ν_τ

- ν interaction enhanced in Earth's crust (producing e , μ , τ , nuclear fragments)
- Only τ can travel long distance through Earth and induce EAS (by decaying after ~ 48 km @ 1 EeV)

Earth-Skimming ν_τ Analysis



No event after $\langle \text{AoP} \rangle$ cut \rightarrow calculate exposure
 \rightarrow flux limit \sim event count limit / exposure