

Dipole anisotropies & local source

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with G.Giacinti, A.Nernov, V.Savchenko, D.Semikoz

Outline of the talk

- 1 **Introduction**
 - ▶ Propagation in magnetic fields
- 2 Dipole anisotropies – calculational approaches
 - ▶ diffusion approach
 - ▶ trajectory approach
- 3 Dipole anisotropy and the transition energy
- 4 Dipole anisotropy in the escape model
 - ▶ Anisotropy from background of all sources
 - ▶ Anisotropy of a single source
 - ▶ single source: other signatures
- 5 Conclusions

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- slope of power spectrum $\mathcal{P}(k) \propto k^{-\alpha}$ determines **energy dependence** of diffusion coefficient $D(E) \propto E^{\beta}$ as $\beta = 2 - \alpha$:

Kolmogorov	$\alpha = 5/3$	\Leftrightarrow	$\beta = 1/3$
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- observed energy spectrum** of primaries:

- ▶ injection: $dN/dE \propto E^{-\alpha}$

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$\alpha = 3/2$ and $\beta = 1/2$ “simplest” combination, degeneracies D/h

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- anisotropy** $\delta_i = -3D_{ij}\nabla_j \ln(n)$

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 - ▶ known for bursting case
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[Lee '72, Blasi & Amato '12]

$$\delta \propto \langle \mathbf{J} \rangle + \langle \delta \mathbf{J} \delta \mathbf{J} \rangle + \dots$$

- ▶ G known for purely turbulent field – what happens for $\mathbf{B}(\mathbf{x}) = \mathbf{B}_{\text{reg}}(\mathbf{x}) + \mathbf{B}_{\text{rms}}(\mathbf{x})$?

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- **both:**
 - ▶ only weak connection of $D_{ij}(\mathbf{x})$ to GMF
 - ▶ diffusion breaks down around the knee

Trajectory approach:

- use **model** for **Galactic magnetic field**
- calculate trajectories $\boldsymbol{x}(t)$ via $\boldsymbol{F}_L = q\boldsymbol{v} \times \boldsymbol{B}$.

[Jansson, Farrar '12]

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- + valid at all energies
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- + includes all info of current GMF models
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- numerically expansive
- 2 methods used:
 - ▶ backward propagation à la Karakula
 - ▶ forward propagation for single sources

[Jansson, Farrar '12]

Anisotropy for Galactic sources

[Giacinti et al. '11]

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- **backtrace** $N = 10^4$ **anti-particles** injected with random direction \hat{r} from Earth

Anisotropy for Galactic sources

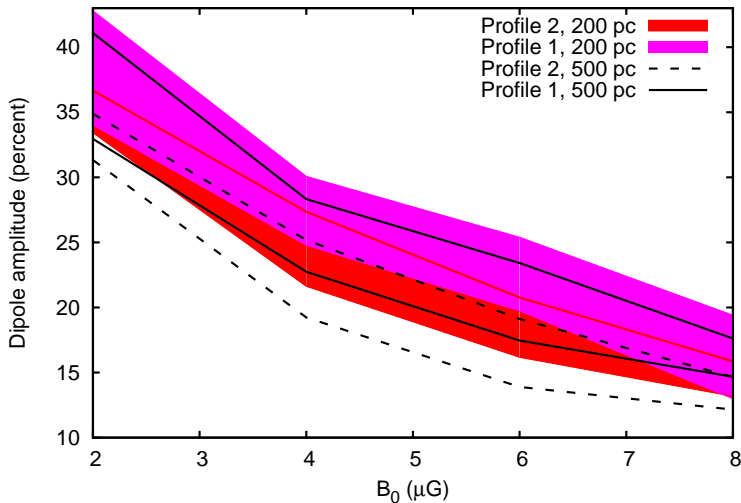
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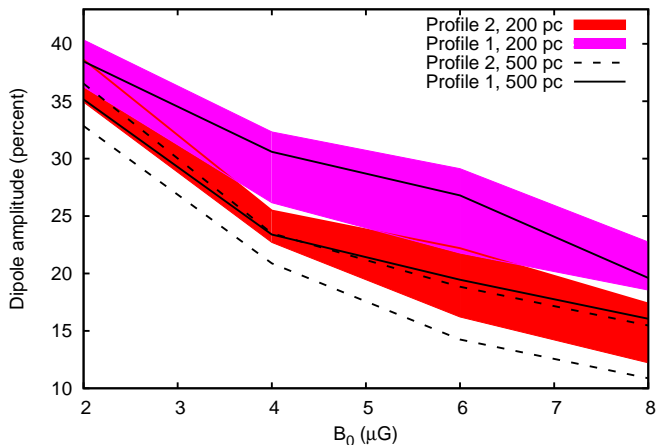
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- backtrace $N = 10^4$ anti-particles injected with random direction \hat{r} from Earth
- weight = path length in the source region
- dipole $d = 3N^{-1} \sum_i w_i \hat{r}_i$

Anisotropy of protons at $E = 10^{18}$ eV – Kolmogorov

- protons excluded for all reasonable parameters

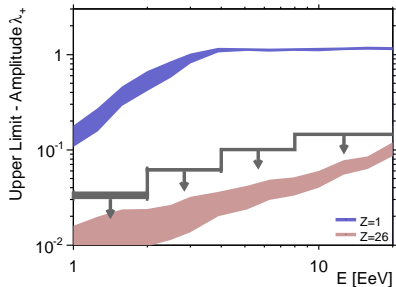
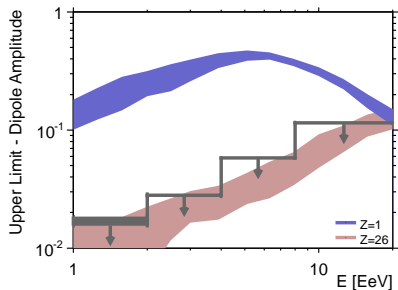
Anisotropy of protons at $E = 10^{18}$ eV – Kraichnan

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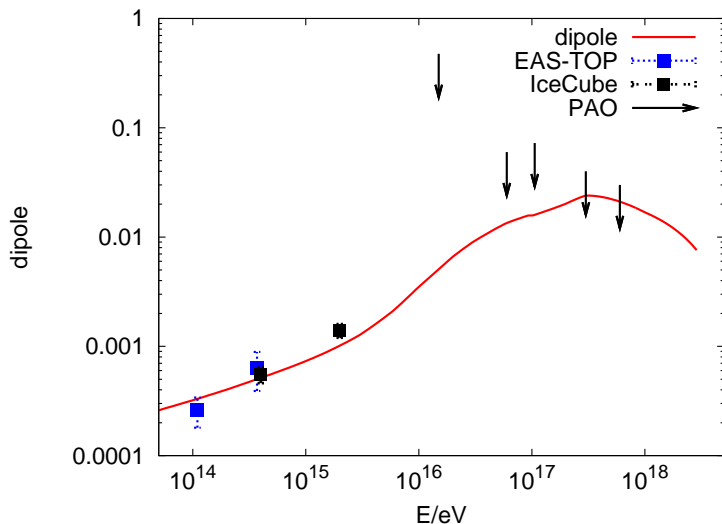
⇒ measuring protons at $E = 10^{18}$ eV means fixing transition energy

Updated PAO results:

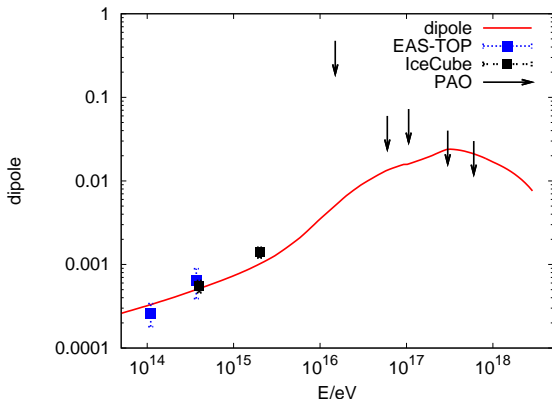
- first 2-dim. analysis
- repeat Giacinti et al. analysis with more statistics:



Knee from CR escape: dipole anisotropy



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- assumes $D(E) \propto 1/X(E)$
- phase changes $10^{15} - 10^{18}$ eV.

Anisotropy of a single source

- if **only turbulent field**:
diffusion = random walk = free quantum particle

- number density is Gaussian with $\sigma^2 = 4DT$

$$\delta = \frac{3D}{c} \frac{\nabla n}{n} = \frac{3R}{2T}$$

- what happens for general fields?

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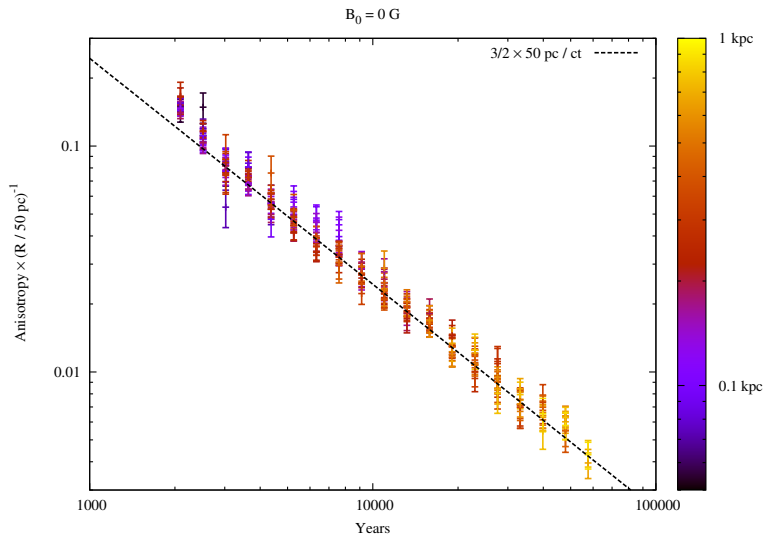
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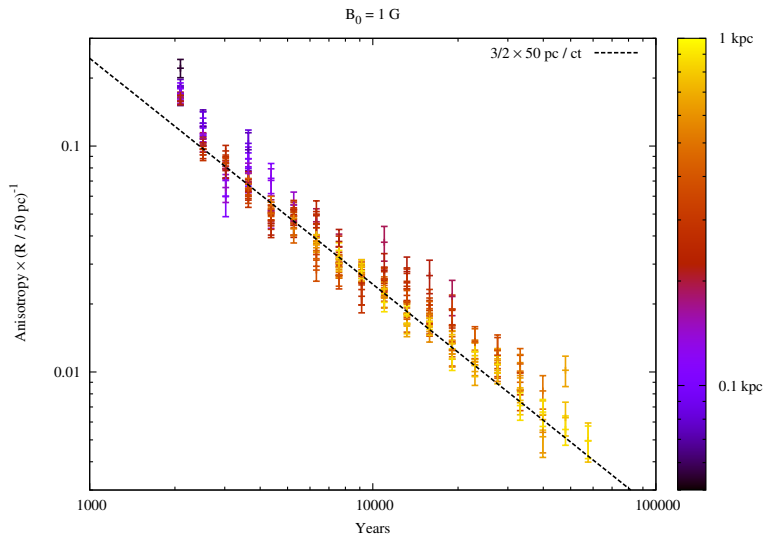
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Anisotropy of a single source: only turbulent field



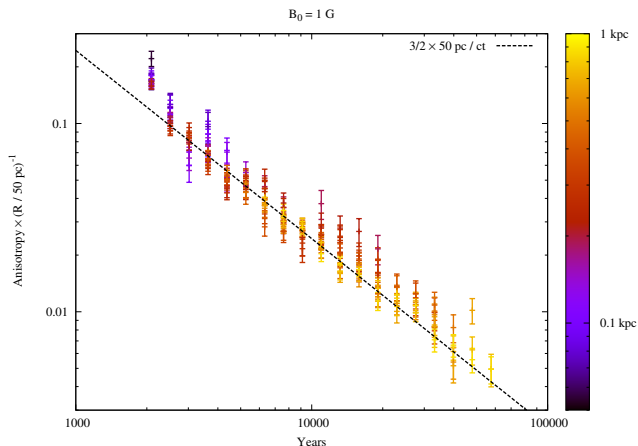
[Savchenko, MK, Semikoz '15]

Anisotropy of a single source: plus regular



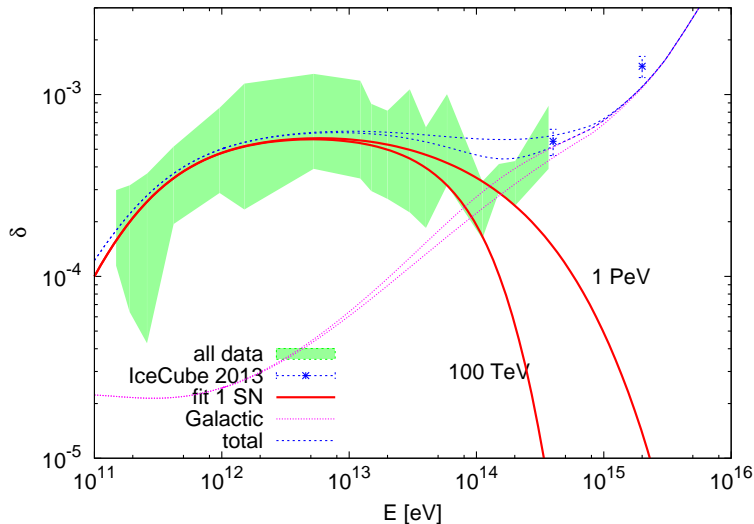
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Anisotropy of a single source:



- regular field changes $n(\boldsymbol{x})$, but keeps it Gaussian
 \Rightarrow no change in δ

Anisotropy of a single source:



[Savchenko, MK, Semikoz '15]

Single source: other signatures

- 2 Myr SN explains anomalous ^{60}Fe sediments

[Ellis+ '96]

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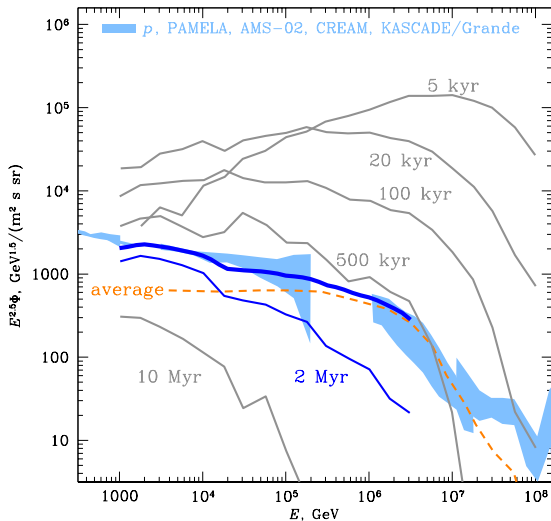
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[Ellis+ '96]

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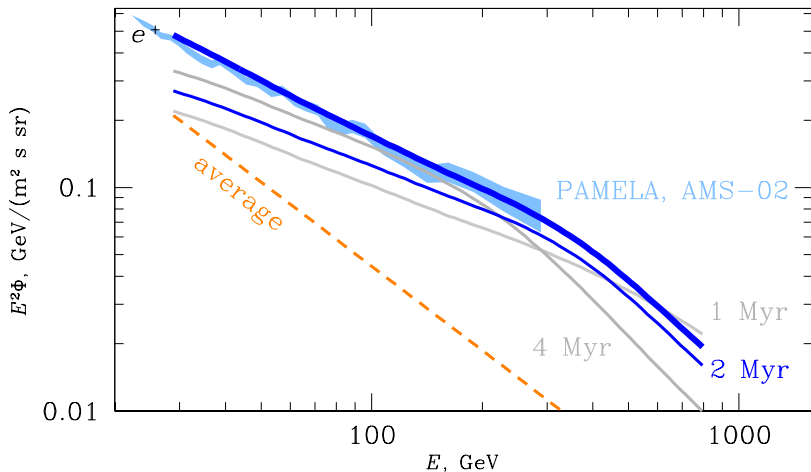
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- may responsible for **different slopes of local p and nuclei fluxes**

Single source: proton flux



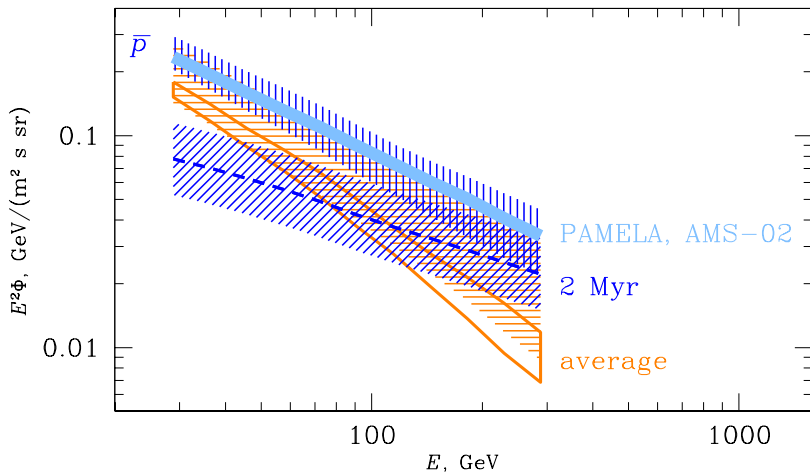
[MK, Neronov, Semikoz '15]

Single source: positrons



[MK, Neronov, Semikoz '15]

Single source: antiprotons



[MK, Neronov, Semikoz '15]

Conclusions

1 Single source: anisotropy

- ▶ dipole formula $\delta = 3R/2T$ holds universally in quasi-gaussian regime
- ▶ plateau of δ points to dominance of single source

2 Single source: antimatter

- ▶ consistent explanation of p , \bar{p} and e^+ fluxes
- ▶ consistent with ^{60}Fe and δ

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